

Water in the terrestrial planet-forming zone of the PDS 70 disk

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Terrestrial and sub-Neptune planets are expected to form in the inner (< 10 AU) regions of protoplanetary disks.¹ Water plays a key role in their formation,^{2–4} although it is yet unclear whether water molecules are formed in-situ or transported from the outer disk.^{5,6} So far *Spitzer Space Telescope* observations have only provided water luminosity upper limits for dust-depleted inner disks,⁷ similar to PDS 70, the first system with direct confirmation of protoplanet presence.^{8,9} Here we report JWST observations of PDS 70, a benchmark target to search for water in a disk hosting a large (~ 54 AU) planet-carved gap separating an inner and outer disk.^{10,11} Our findings show water in the inner disk of PDS 70. This implies that potential terrestrial planets forming therein have access to a water reservoir. The column densities of water vapour suggest in-situ formation via a reaction sequence involving O, H₂, and/or OH, and survival through water self-shielding.⁵ This is also supported by the presence of CO₂ emission, another molecule sensitive to UV photodissociation. Dust shielding, and replenishment of both gas and small dust from the outer disk, may also play

a role in sustaining the water reservoir.¹² Our observations also reveal a strong variability of the mid-infrared spectral energy distribution, pointing to a change of inner disk geometry.

JWST-MIRI^{13,14} Medium Resolution Spectrometer¹⁵ (MRS) ($\lambda/\Delta\lambda \sim 1600 - 3400$) observations of PDS 70 were taken as part of the guaranteed time MIRI mid-INfrared Disk Survey (MINDS; see Methods and Extended Data Fig. 1). The complete spectrum of PDS 70 shows several distinct traits (Fig. 1), which stands out with respect to other T Tauri disks.^{16,17}

A significant flux offset - up to a factor of 1.5 at wavelengths $\geq 18 \mu\text{m}$ - is found between the MIRI and the archival *Spitzer*-IRS low-resolution ($\lambda/\Delta\lambda \sim 60 - 100$) spectra recorded with 15 years and one day time difference. This discrepancy is too large to be explained by calibration uncertainties; the absolute uncertainty for both IRS and MIRI is $\sim 5\%$ for the $4.9 - 22.5 \mu\text{m}$ range.¹⁸ Similarly, the difference in aperture size of the two spectrographs cannot account for such an offset. Hence, with the current MIRI data reduction, time variability is the most likely explanation for the observed flux differences.

Variability in the mid-IR observed with *Spitzer*-IRS has been mainly attributed to short-wavelength stellar irradiation or to dynamical changes in the inner disk geometry due to the presence of planets.^{19,20} In the case of PDS 70, stellar irradiation is excluded as it would cause an overall increase or decrease in flux contrary to what is observed with WISE Time-series observations (Extended Data Fig. 6). PDS 70 is known to be in a late stage of accretion - with an estimated waning accretion rate^{21,22} of $\sim 10^{-10} M_{\odot} \text{yr}^{-1}$ making it unlikely to explain the significant flux difference. Changes in the scale height of the inner disk wall emitting at shorter wavelengths ($\sim 2 - 8 \mu\text{m}$) can be responsible for shadowing the disk material located further out, resulting in less emission at wavelengths beyond $18 - 20 \mu\text{m}$ (also referred to as "seesaw-like" variability¹⁹). However, for PDS 70 a complete seesaw-like profile is not observed as there is no corresponding increase in flux at the shorter wavelengths. Time variability is also supported by the aforementioned WISE observations which indicate that such variability occurs on short timescales ($\leq 1 \text{yr}$) and that it may indeed be attributed to occulting material located close to the star ($\sim 1 \text{AU}$).

The MIRI spectrum of PDS 70 clearly shows the presence of silicate dust grains that have undergone significant thermal processing (Fig. 1). We attribute the crystalline dust features to enstatite at $9.40 \mu\text{m}$ and forsterite at $11.30 \mu\text{m}$ and $16.40 \mu\text{m}$. The observed dust continuum is well reproduced with a three-component disk model, with a $400 - 600 \text{K}$ surface layer accounting for the bulk of the observed emission (Fig. 2).

The MIRI spectrum also reveals a wealth of water lines, particularly in the $7 \mu\text{m}$ spectral window (Fig. 3). This is the first detection of water in the terrestrial region of a disk already hosting two or more protoplanets. As such, it also provides important clues to theories on the origin of water during terrestrial planet formation in the solar system.^{23,24} We focus on the ro-vibrational transitions of the bending mode of para- and ortho-water in the $7 \mu\text{m}$ region where the brightest lines are observed and contamination by the stellar atmosphere is negligible (Extended Data Fig. 3). This includes strong water blends dominated by lines with excitation $E_u \simeq 2400 - 3200 \text{K}$. Weaker lines are also detected at the 1mJy level, some of them corresponding to more excited levels up to $E_u \simeq 4300 \text{K}$.

Further insight into the origin of water emission is obtained from 0D slab modeling which also has been used to interpret *Spitzer* spectroscopic data.²⁵ The synthetic spectrum of water is calculated from a plane-parallel slab model, where the level populations are in Local Thermodynamic Equilibrium (LTE) at a single excitation temperature T . The other fitting parameters are the line-of-sight column density N within an effective emitting area πR^2 given by its radius R , and the intrinsic line broadening assumed to be $\sigma = 2 \text{km/s}$.²⁵ Note that R does not need to correspond to a disk radius, but could also represent an annulus with the same area or an emission spot breaking the axisymmetry. The best-fit model is then obtained by minimizing the reduced χ^2 between measured and model line fluxes over the individual spectral window around each H₂O line (Extended Data Fig. 4).

The observed H₂O spectrum in the 6.78–7.36 μm spectral region is best fitted with a slab of gas at $T = 600$ K, with an emitting area of radius $R = 0.047$ AU, and a column density of $N = 1.4 \times 10^{18} \text{ cm}^{-2}$. The temperature is mostly determined by the ratio between lines of different E_u , for example, the series of lines in the 7.3 μm region. The column density is set by the ratio between the weaker lines and indicates that the brightest lines are optically thick. The emitting area is constrained by matching the fluxes of the optically thick lines, and points toward a compact emission region. This is further supported by the fact that the detected lines are broad ($\Delta\lambda \sim 0.01 - 0.05$ μm); if the line broadening is caused by the gas kinematics, the FWHM of the line would be about 100 km/s, corresponding to a Keplerian radius of 0.1 AU, consistent with the emitting area deduced from our fit after correction for disk inclination $i = 51.7 \pm 0.1^\circ$.¹¹ Interestingly, we find that our best-fit LTE model of the water emission in the 7 μm region reproduces reasonably well water rotational lines at 15 μm, suggesting that all water emission in the MIRI spectral range originates from inside ~ 0.05 AU under LTE conditions (Extended Data Fig. 5).

Besides water in the 7 μm, 15 μm, and 17 μm regions, other species have been identified whose analysis is postponed to a future study. **CO₂**: The fundamental Q-branch of CO₂ corresponding to the ν_5 bending mode is detected at 14.96 μm (Extended Data Fig. 5). Interestingly, the width of this feature is sensitive to the temperature and indicates cooler gas of $T \simeq 200$ K in the optically thin regime. **H₂**: The pure rotational H₂ $S(5)$ and H₂ $S(1)$ lines are detected at 6.91 μm and 17 μm (Fig. 1). We note that the H₂ $S(2)$, $S(3)$, and $S(4)$ lines coincide with the broad silicate emission feature and thus establishing their presence needs to await an in-depth analysis of this dust feature.

Spitzer-IRS observations detected water in $\sim 50\%$ of dust-rich inner disks around T Tauri stars²⁶, but obtained only upper limits for disks with large inner dust gaps or cavities defined by a mid-IR spectral index $n_{13-30} > 0.9$ (Fig. 4).⁷ The detection of water vapour in the PDS 70 MIRI spectrum demonstrates that PDS 70 has maintained to some degree the physical and chemical conditions of dust-rich inner disks in its terrestrial planet-forming zone despite the presence of a remarkably large gap (see "Origin of water in PDS 70" in Methods). Our LTE slab model only provides a first quantitative analysis of the H₂O emission. Non-LTE effects could lead to sub-thermal line emission which would make our estimated emitting area a lower limit. In T Tauri disks with strong radial temperature gradients, the water lines are expected to originate from different regions of the disk depending on their upper energy level and Einstein-A coefficients.²⁷ Detailed modelling using a realistic disk structure and including non-LTE effects such as IR radiative pumping is needed in the future to further constrain the distribution of H₂O across the inner disk. However, this first analysis already proves that the inner disk of PDS 70 is rich in water and the inferred slab model parameters are roughly consistent with a detailed thermo-chemical model.²⁸

The luminosity of the 17-μm water lines is two orders of magnitude weaker for PDS 70 than DoAr 44.^{7,29} DoAr 44 is a system with similar properties to PDS 70, but characterized by $n_{13-30} < 0.9$ (Fig. 4). This result points to a colder water reservoir in PDS 70, and is consistent with the lower luminosity and lower accretion rate of PDS 70.^{21,30} This work opens a new window on the origin of water in protoplanetary disks by showing that MIRI-MRS can now detect very weak ($\lesssim 5$ mJy) water lines in the innermost regions of disks with large gaps, and hence that the presence of water in the terrestrial planet-forming zone of dust-depleted inner disks is not as rare as previously thought.

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Author contributions

G.P. and V.C. performed the data reduction, supported by D.G., M.S., I.A., and J.B. G.P., I.K., V.C., B.T., L.B.F.M.W., G.O. and T.H. wrote the manuscript. G.P., B.T. and S.L.G. did the line analysis. L.B.F.M.W. carried out the dust continuum analysis. G.O. performed the correction for photospheric emission. J.B. performed the reduction of the *Spitzer* dataset. T.H. and I.K. planned and co-led the MINDS guaranteed time program. All authors participated in either the development and testing of the MIRI instrument and its data reduction, in the discussion of the results, and/or commented on the manuscript.

Competing Interests

The authors declare no competing financial interests.

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Data Availability

The original data analysed in this work are part of the Guaranteed Time Observation (GTO) program 1282 (PI: Th. Henning) with number 66 and will become public on August 1, 2023 on the MAST database (<https://mast.stsci.edu>). The portion of the spectrum presented in Fig. 3 is available on Zenodo at <https://zenodo.org/record/7991022>. The spectroscopic data for water can be downloaded

from the HITRAN database (<https://hitran.org>). The Spitzer-IRS spectrum plotted in Fig. 1 is part of the Spitzer-IRS GTO program 40679 (PI: G. Rieke). The spectrum was extracted and calibrated using private codes^{56,64} and is available on Zenodo at <https://zenodo.org/record/7991022>. The optical constants of the dust species considered in the fitting procedure for the dust continuum can be downloaded from the HJPDOC database (<https://www2.mpia-hd.mpg.de/HJPDOC>).

Code Availability

The slab model used in this work is a private code developed by B.T. and collaborators. It can be obtained from B.T. upon request. The synthetic spectra presented in this work can be reproduced using the slabspec code, which can be found at <https://doi.org/10.5281/zenodo.4037306>. The fitting procedure for the dust continuum uses the publicly available MultiNest Bayesian fitting algorithm (<https://github.com/JohannesBuchner/MultiNest>) and the PyMultiNest package (<https://github.com/JohannesBuchner/PyMultiNest>). Figures were made with Matplotlib version 3.5.1. under the Matplotlib license at <https://matplotlib.org/>.

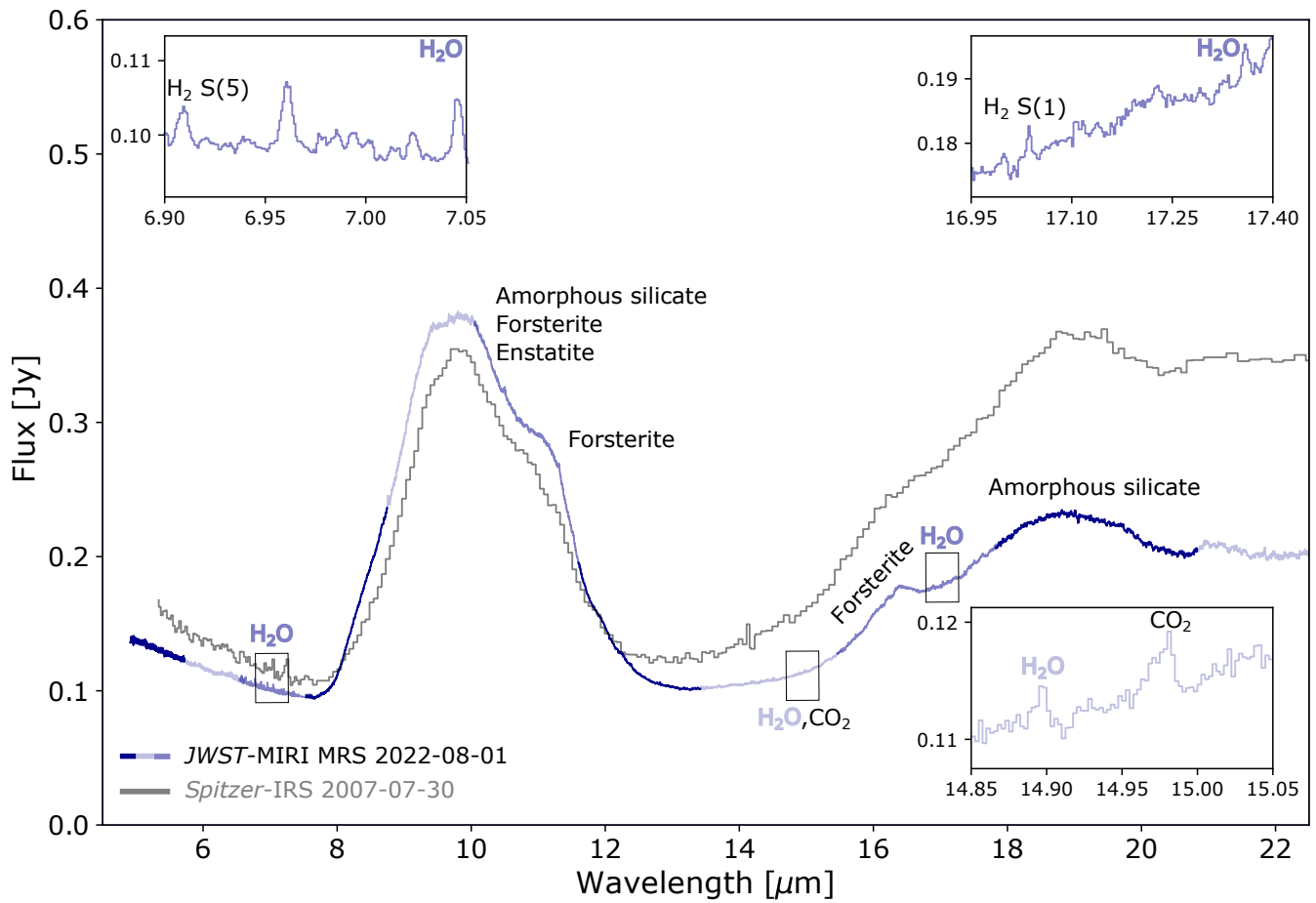


Figure 1. JWST-MIRI-MRS spectrum of PDS 70. The spectrum is a composite of the colour-coded short, medium, and long sub-bands of the four MIRI-MRS Integral Field Units (IFUs).¹³ The *Spitzer*-IRS spectrum of PDS 70 is also shown in grey. The major dust features are labelled. The spectrum is dominated by exceptionally prominent silicate emission at 10 and 18 μm and it clearly shows a number of crystalline dust features. The much higher sensitivity and spectral resolution of MIRI-MRS compared with previous *Spitzer* spectra allows us to detect for the first time an inner disk gas reservoir by showing weak emission of water vapour and carbon dioxide as well as two molecular hydrogen lines. The insets show the ro-vibrational and rotational transitions of ortho- and para- H_2O , the H_2 S(1) and S(5) rotational lines, and the ν_5 bending mode of CO_2 .

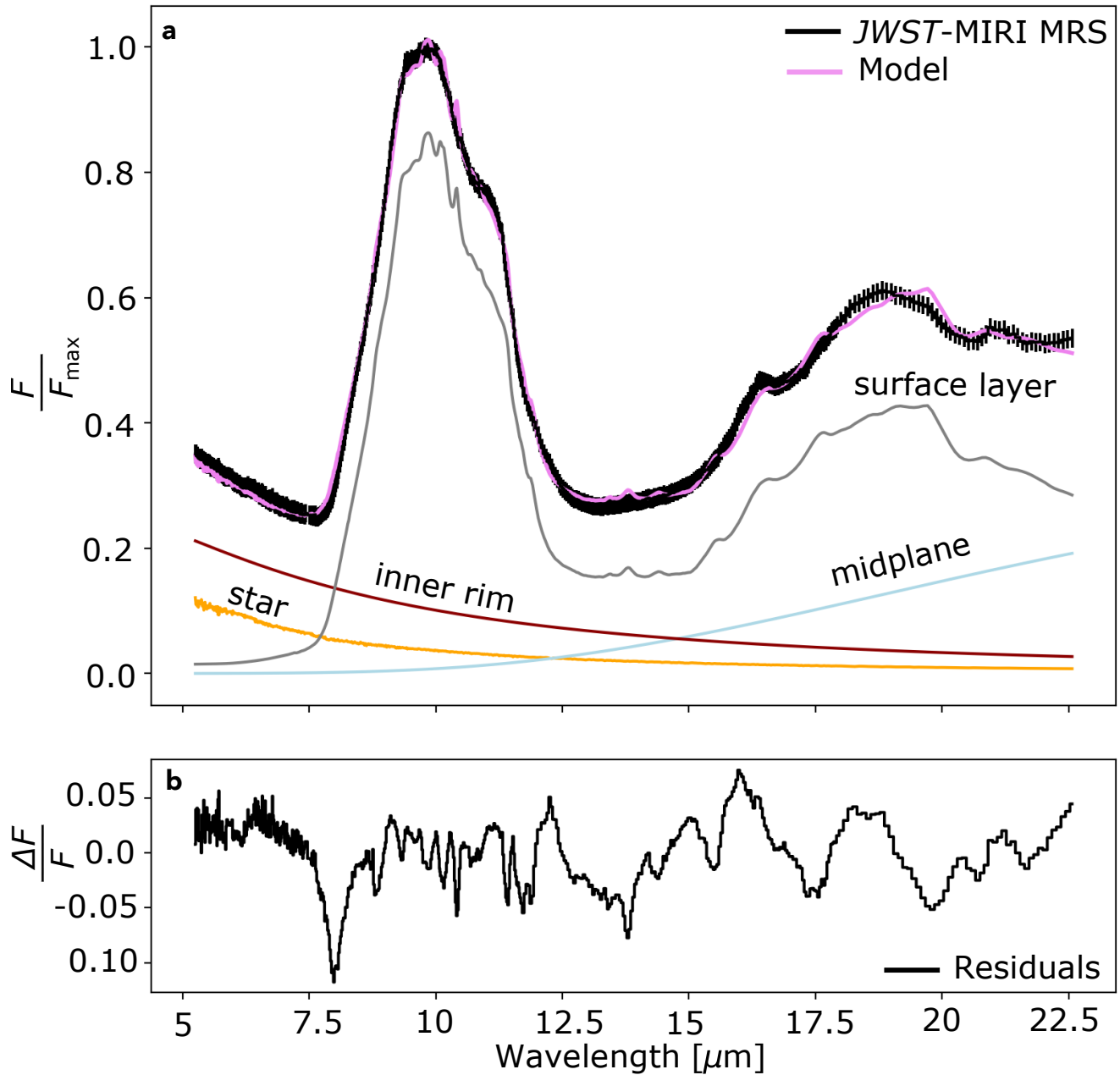


Figure 2. Dust continuum fit to the MIRI spectrum of PDS 70. **a**, The disk model has three spectral components: an inner rim, an optically thick mid-plane disk layer, and an optically thin warm disk surface layer. The stellar photospheric emission is represented by a stellar atmosphere model (see Methods for further details). The surface layer component dominates the MIRI spectrum in the 8–22.5 μm wavelength range. Its temperature is constrained to be between 400 and 600 K. The silicate emission at 8–12 μm is consistent with a population of optically thin dust grains with typical sizes of 0.1–2 μm . A significant contribution from an optically thick dust component is excluded because of the high silicate peak/continuum ratio of $\simeq 4$.³¹ **b**, The residuals on the dust continuum fit.

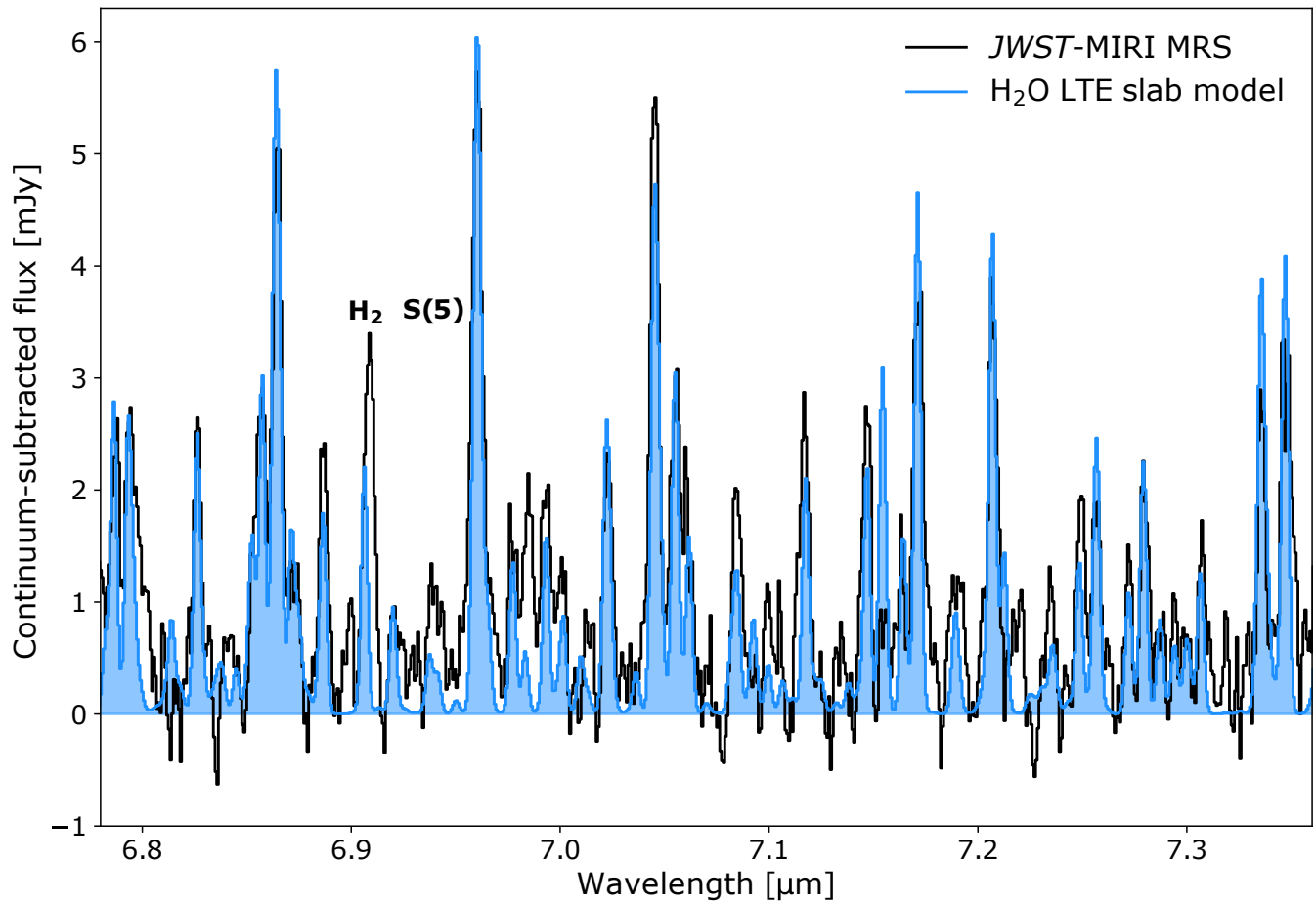


Figure 3. Continuum-subtracted spectrum showing H₂O emission in the 7 μm region and the best-fit LTE slab model. The best-fit model (blue) has $T = 600$ K, $N(\text{H}_2\text{O}) = 1.4 \times 10^{18} \text{ cm}^{-2}$, and $R = 0.047$ AU. The H₂ S(5) line is labelled on top of the spectrum.

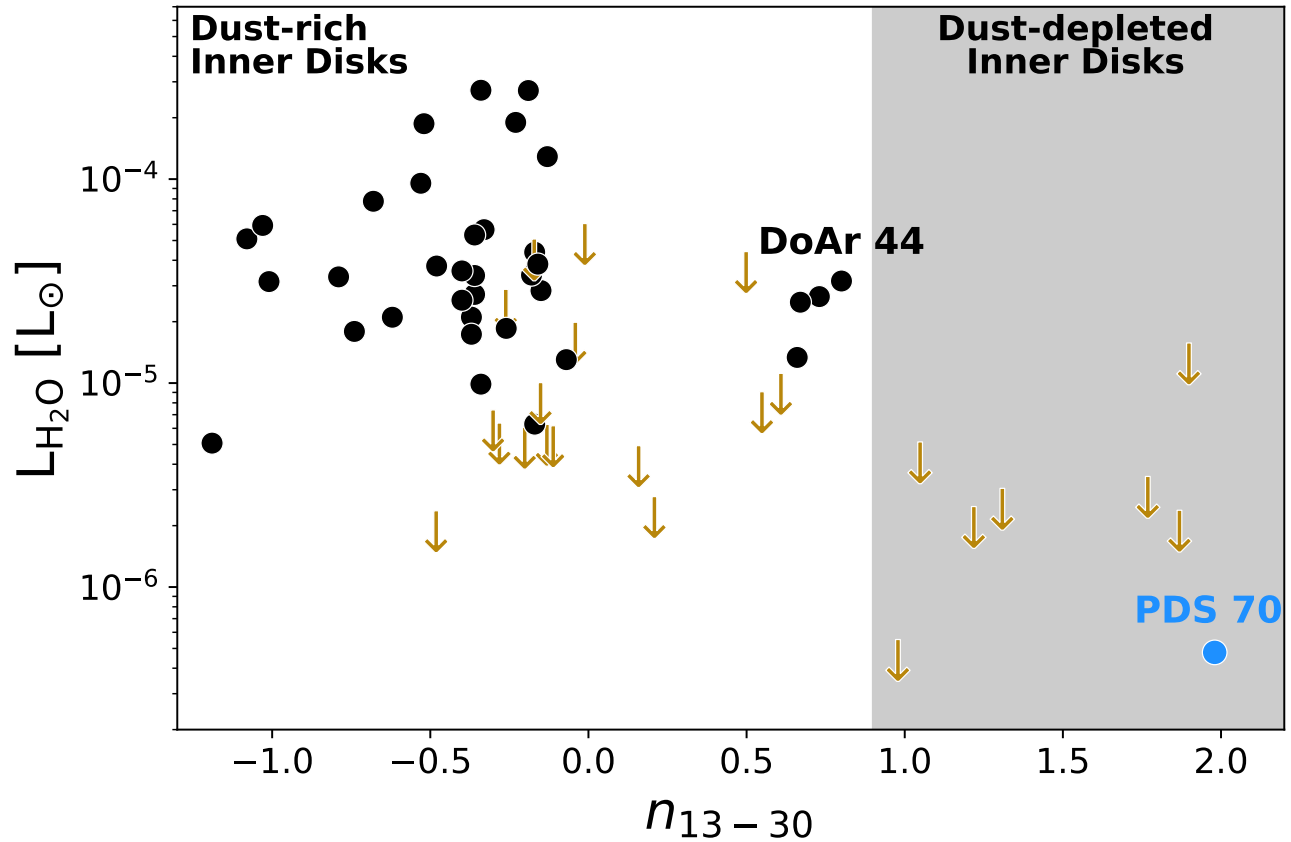


Figure 4. Comparison between water luminosity and mid-infrared spectral index (n_{13-30}) for a sample⁷ of protoplanetary disks. n_{13-30} is a diagnostic of the presence and size of inner disk dust cavities: $0.9 < n_{13-30} < 2.2$ corresponds to disks with large gaps and/or cavities.^{32,33} Black dots represent disks with mid-infrared water detections. Disks for which only upper limits were obtained are shown as gold arrows. The grey shaded area highlights the location of dust-depleted inner disks with PDS 70 shown as a blue dot. The water line flux used to compute the water luminosity of PDS 70 is calculated as described in a previous work.⁷ The *Spitzer* spectrum is used to estimate n_{13-30} for PDS 70 to be consistent with the other targets. *Spitzer*-IRS obtained only water luminosity upper limits for disks characterized by n_{13-30} greater than 0.9. Below $10 \mu\text{m}$, IRS provided only a spectral resolution of $R \sim 100$, preventing a comprehensive view of water in the innermost regions.³⁴ DoAr 44 is a system schematically similar to PDS 70.²⁹ The two stars have comparable age and spectral types K3 and K7 respectively; DoAr 44 has a higher mass accretion rate of $\dot{M}_{\text{acc}} \sim 10^{-8} M_{\odot} \text{ yr}^{-1}$.³⁰ The cavity size of DoAr 44 is 34 AU^{7,29}, smaller than that of PDS 70 ($\sim 54 \text{ AU}$ ¹¹). Both systems have small inner disks based on VLTI-GRAVITY³⁵, VLT-SPHERE and ALMA data^{8,11}, and in both systems, the water emission is contained to within 1 AU.²⁹ The water luminosity of PDS 70 is two orders of magnitude weaker than that of DoAr 44, pointing to a colder water reservoir in PDS 70. This is consistent with the lower luminosity and lower accretion rate of PDS 70.

Methods

PDS 70 system. PDS 70 (V1032 Cen) is a K7-type star in the Upper Centaurus-Lupus subgroup ($d=113.4\pm 0.5$ pc³⁶) in a late stage of accretion²² with an estimated age of 5.4 ± 1.0 Myr.³⁷ The disk around PDS 70^{38–40} hosts two actively accreting protoplanets: PDS 70 b and PDS 70 c which reside in a ~ 54 AU-annular gap between an inner and outer disk.^{8,9} The presence of an inner dusty disk in the PDS 70 system has been inferred from both near-IR scattered light and ALMA images.^{10–12} The $855\ \mu\text{m}$ dust continuum emission from the innermost disk regions is confined within the orbit of PDS 70 b (~ 22 AU; Extended Data Fig. 1), putting an upper limit to the inner disk radial extent of ~ 18 AU.¹² A population of small dust grains may be responsible for the observed inner disk emission although the current dust mass estimates, inferred from the low near-IR excess in the spectral energy distribution, prevent to exclude the simultaneous presence of small and large dust grains.⁴¹

Observations and data reduction. The PDS 70 disk (CD-40-8434) was observed with the Mid-InfraRed Instrument (MIRI)^{13,14} mounted on JWST on August 1, 2022 as part of the Guaranteed Time Observation (GTO) program 1282 (PI: Th. Henning) with number 66. The disk was observed in FASTR1 readout mode with a 4-point dither pattern in the negative direction for a total on-source exposure time of 4,132 s. The Medium Resolution Spectroscopy (MRS)¹⁵ mode was used, which has four Integral Field Units (IFUs). Each IFU (referred to as channel) covers a different wavelength range and splits the field of view into spatial slices. Calibration and processing of IFU observations produces 3-dimensional spectral cubes. The latter are used to extract a final spectrum covering the MIRI $4.9\text{--}22.5\ \mu\text{m}$ range and is a composite of the four IFUs: channel 1 ($4.9\text{--}7.65\ \mu\text{m}$; $R\sim 3400$), channel 2 ($7.51\text{--}11.71\ \mu\text{m}$; $R\sim 3000$), channel 3 ($11.55\text{--}18.02\ \mu\text{m}$; $R\sim 2400$), and channel 4 ($17.71\text{--}22.5\ \mu\text{m}$; $R\sim 1600$). Each channel is in turn composed by three sub-bands: SHORT (A), MEDIUM (B), and LONG (C) leading to a total of twelve wavelength bands.

We processed the PDS 70 data using a hybrid data reduction pipeline made from the combination of the JWST Science Calibration pipeline⁴² (v1.8.4) stages 1 to 3, with dedicated routines based on the VIP package^{43,44} for bad pixel correction, background subtraction and removal of spikes affecting the final spectrum. Specifically, data reduction proceeded as follow: (i) the class `Detector1` of the JWST pipeline was used to process uncalibrated raw data files using CRDS context `jwst_1019.pmap` and default parameters; (ii) apart from pixels flagged in the DQ extension, we identified additional bad pixels with both an iterative sigma clipping algorithm and through a cross-shaped match filter, and corrected them using a 2D Gaussian kernel; (iii) `Spec2` was then used with default parameters, but the background subtraction was skipped, and dedicated reference files⁴⁵ for photometric and fringe flat calibrations were adopted; (iv) as no dedicated background observation was taken, we leveraged the four-point dither pattern to obtain a first guess on the background map, then refined it using a median-filter which both smoothed the background estimate and removed residual star signals from it; (v) `Spec3` was then run with default parameters, apart from the `master_background` and `outlier_detection` steps which were turned off, in the latter case to avoid spurious spectral features resulting from under-sampling of the PSF; (vi) we recentered the spectral cubes by applying the shifts maximizing the cross-correlation between cube frames, and found the location of the PSF centroid with a 2D Gaussian fit on the median image of each aligned cube; (vii) spectra were then extracted with aperture photometry in 2.5 FWHM apertures centered on the centroid location (with the FWHM equal to $\sim 1.22\ \lambda/D$), corrected for both aperture size using correction factors¹⁸, and spikes affecting individual spaxels included in the aperture; (viii) spectra were finally corrected for residual fringes at the spectrum level and the bands were stitched together based on the level of the shorter wavelength bands (these rescaling factors were systematically within 3% of the photometric solution). Spurious data reduction artefacts were masked at $5.12\ \mu\text{m}$,

5.90 μm , 7.45 μm , 7.50 μm . The uncertainty associated to each photometric measurement considers both Poisson and background noise, combined in quadrature. The former is an output from the JWST pipeline, while for the latter we propagated our background estimate obtained in step (iv) through `spec3`, and considered the standard deviation of the fluxes inferred in independent 2.5 FWHM apertures as a proxy for the background noise uncertainty. The final relative uncertainties range from $\sim 0.1\%$ to $\sim 1.1\%$ with respect to the continuum at the shortest and longest wavelengths considered in this work (4.9 and 22.5 μm , respectively).

Local continuum fit. Extended Data Fig. 2 shows the local baseline fit for the 7 μm region. The continuum level is determined by selecting line-free regions and adopting a cubic spline interpolation (`scipy.interpolate.interpld`). This continuum is then subtracted from the original data to produce the spectrum shown in Fig. 3.

Correction for the photospheric emission. The observed near-IR colour index, $J - K_s = 1.01$ (2MASS), indicates a small colour excess, $E(J - K_s) \approx 0.16$ which could be due to either interstellar extinction or a true excess in the K_s band, or a combination of the two. By assuming that the brightness in the K_s band is essentially due to photospheric emission we can, by using a model atmosphere kindly provided by P. Hauschildt (priv. communication), extrapolate the contribution into the mid-IR spectral region. The parameters used for the model atmosphere of the PDS 70 K7-star are $T_{\text{eff}} = 4000$ K, $\log(g) = 4.5$, and solar metallicity. At 5 μm the photospheric contribution amounts to 56 mJy, i.e., 44% of the observed flux density. At longer wavelengths, in the 7 μm region where the water emission is detected, the photosphere amounts to a third of the observed flux density and the subtraction of the photospheric contribution just marginally alters the continuum-subtracted spectrum (Extended Data Fig. 3).

Slab models fits. The molecular lines are analysed using a slab approach that takes into account optical depth effects. The level populations are assumed to be in Local Thermodynamical Equilibrium (LTE) and the line profile function to be Gaussian with a full-width half maximum of $\Delta V = 4.7$ km s $^{-1}$ ($\sigma = 2$ km/s).²⁵ The line emission is assumed to originate from a slab of gas with a temperature T and a line of sight column density of N . Under these assumptions and neglecting mutual line opacity overlap, the frequency-integrated intensity of a line (in erg s $^{-1}$ cm $^{-2}$ sr $^{-1}$) is computed following:

$$I = \frac{\Delta V}{2\sqrt{\ln 2}\lambda_0} B_{\nu_0}(T) \int_{-\infty}^{+\infty} (1 - \exp(-\tau_0 e^{-y^2})) dy, \quad (1)$$

where $B_{\nu_0}(T)$ is the Planck function, λ_0 is the rest wavelength of the line, and τ_0 is the optical depth at the line center ν_0 , with:

$$\tau_0 = \sqrt{\frac{\ln 2}{\pi}} \frac{A_{ul} N \lambda_0^3}{4\pi \Delta V} \left(x_l \frac{g_u}{g_l} - x_u \right). \quad (2)$$

In this equation, x_l and x_u denote the level population of the lower and upper states, and g_l and g_u their respective statistical weights, and A_{ul} the spontaneous downward rate of the transition. The line intensity is then converted into integrated flux F_{ν_0} (in erg s $^{-1}$ cm $^{-2}$) assuming an effective emitting area of πR^2 and a distance to the source d as:

$$F_{\nu_0} = \pi \left(\frac{R}{d} \right)^2 I. \quad (3)$$

We note that neglecting mutual line overlap for H $_2$ O when calculating the line intensity is a valid approximation for $N(\text{H}_2\text{O}) \lesssim 10^{20}$ cm $^{-2}$ and significantly reduces the computational time.⁴⁶ Finally,

the spectrum is convolved and sampled in the same way as the observed spectrum⁴⁷ and all lines are then summed to prepare a total synthetic spectrum. The molecular data, i.e., line positions, Einstein A coefficients and statistical weights stem from a previous work.⁴⁸

Fitting procedure for H₂O vapour lines. The LTE slab model described above is then used to fit the H₂O lines in the 6.78–7.36 μm region following a χ^2 method. First, an extended grid of models is computed varying the total column density from 10^{15} to 10^{20} cm^{-2} in steps of 0.17 in log10-space and the temperature from 100–1400 K in steps of 50 K. We further assume an ortho-to-para ratio of 3. For each set of free parameters (N, T, R), a synthetic spectrum is calculated at the spectral resolution $\Delta\lambda/\lambda = 2000$ and rebinned to the spectral sampling of the observed spectrum using the `slabspec` python code.⁴⁹ The adopted spectral resolution is lower than the nominal MIRI-MRS spectral resolution in channel 1;¹⁵ it was selected to account for the observed line broadening. This spectrum is further used to compute the χ^2 value on a spectral channel basis. Specific spectral windows are chosen to avoid contamination by other gas features. This includes all spectral channels falling within 0.02 μm (1000 km/s) of any HI lines and a 0.01 μm wide spectral window at the position of the S(5) line of H₂ at 6.91 μm . In order to mitigate the errors induced by the continuum subtraction procedure, we also include only spectral elements falling within 0.004 μm of a water line. For each value of (N, T), the χ^2 is then minimized by varying the emitting size πR^2 . The resulting χ^2 map is shown in Extended Data Fig. 4 together with the best fit emitting radius. The confidence intervals are estimated following a previous work⁵⁰ and using a representative noise level of $\sigma = 0.15$ mJy. We note that due to the large number of lines in this crowded region, there is little space to determine the noise on the continuum. Therefore, we estimate the noise level between 7.72 μm and 7.73 μm to avoid contamination by H₂O and HI lines.

Origin of water in PDS 70. At the typical densities of inner disk regions ($n_{\text{H}} \geq 10^8$ cm^{-3}) the chemistry can rapidly reach steady-state conditions and water vapour can form from a simple reaction sequence involving O, H₂, and OH. Water and OH absorb efficiently in the UV (i.e., water and OH shielding), ensuring the survival of water molecules even in regions of reduced dust opacity.⁵ This mechanism by itself is able to account for the column densities of water vapour detected in this work and is supported by the presence of CO₂ emission. Small grains in the inner disk provide additional UV shielding. One question that naturally arises is whether the water vapour in PDS 70 formed prior to the formation of the giant protoplanets within the gap or whether there is a continuous supply of gas from the outer to the inner disk regions. ALMA high spatial resolution CO observations reveal the presence of gas inside the gap.^{11,51} Observations and models find the gap to be gas depleted¹¹ (2-3 orders of magnitude assuming an r^{-1} surface density profile) and dust depleted⁵², but not empty. One possibility could be that a population of water-containing dust particles is able to filter through the orbits of PDS 70 b and PDS 70 c, enriching the inner disk reservoir.¹² Experimental evidence suggests that water chemically bound to complex silicates can be preserved to temperatures up to 400 – 500 K^{53,54} and thus survive in the regions probed by our observations inside the water snowline. We note that some degree of dust filtering is expected with gas replenishment, as small dust particles can couple to the gas. Therefore a replenishment of both gas and dust from the outer disk to sustain the water reservoir and hence PDS 70’s accretion rate is possible.

Fitting procedure for the dust continuum. The 4.9 – 22.5 μm dust continuum is analysed using a two-layer disk model for the dust emission.⁵⁵ This model was successfully applied to *Spitzer*-IRS spectra of planet forming disks;⁵⁶ we follow the same modeling approach here. We rebin the spectrum by averaging 15 spectral points and assign errors σ to the rebinned spectral points assuming a normal error distribution with equal weights for each individual spectral element. The stellar photospheric emission is represented by a stellar atmosphere model fitted to optical and near-infrared photometry. The disk model has three spectral components: (1) a hot inner disk F_{rim} , (2) an optically thick mid-plane disk layer F_{mp} , and (3) an

optically thin warm disk surface layer F_{sur} . The dust grains representing the disk components are assumed to have power-law temperature distributions, and each is characterized by a minimum and a maximum temperature T_{atm} . The disk surface layer is assumed to consist of a number of dust species i with different chemical compositions, and with a fixed number of grain sizes j , all emitting at the same temperatures. The total disk flux can then be written as:

$$F_{\text{v}} = F_{\text{v,rim}} + F_{\text{v,mp}} + F_{\text{v,sur}} \quad (4)$$

where

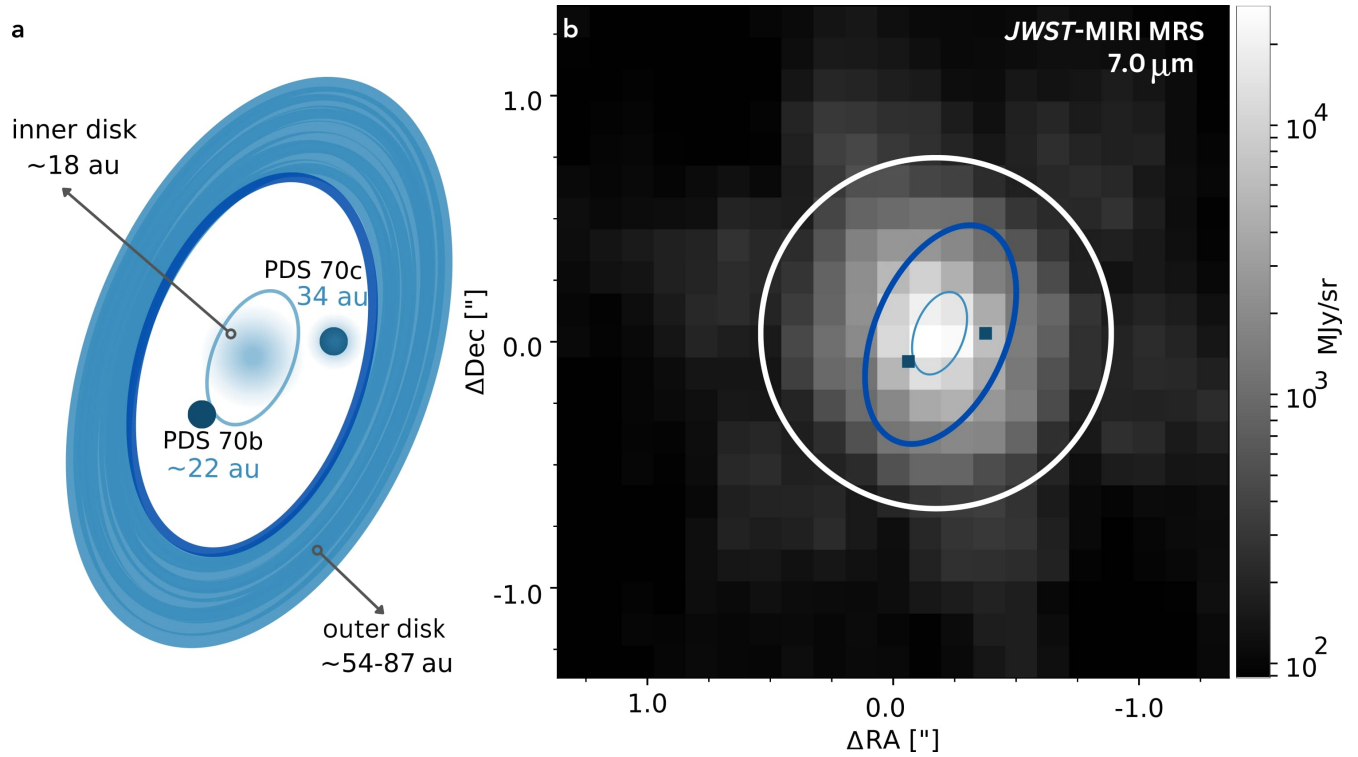
$$F_{\text{v,sur}} = \sum_{i=1}^n \sum_{j=1}^m D_{i,j} \kappa_{i,j} \int_{T_{\text{atm,max}}}^{T_{\text{atm,min}}} \frac{2\pi}{d^2} B_{\text{v}}(T) T^{\frac{2-q_{\text{atm}}}{q_{\text{atm}}}} dT \quad (5)$$

and $B_{\text{v}}(T)$ is the Planck function, q_{atm} is the power law exponent for the temperature gradient in the disk surface layer, $\kappa_{i,j}$ are the opacities in cm^2g^{-1} of dust species i with grain size j , d is the distance to the star, and $D_{i,j}$ are normalization factors.⁵⁵ We use three grain compositions (with SiO_2 , SiO_3 , and SiO_4 stoichiometry)^{57–61} and both amorphous and crystalline lattice structures to capture the rich spectral structure evident in the MIRI data. The choice of this set of compositions is based on previous analyses,⁵⁶ that showed that this set of materials is able to capture most spectral variations in planet-forming disks observed with *Spitzer*-IRS. We use either 2 or 3 grain sizes (i.e., $0.1 \mu\text{m}$, $2 \mu\text{m}$, $5 \mu\text{m}$) for each of the dust species. In total the model has 23 fitting parameters. We use the MultiNest Bayesian fitting algorithm⁶² and the PyMultiNest package⁶³ to find the best-fit parameters. The resulting fit and the separate spectral components (star, inner rim, midplane, surface layer) are shown in Fig. 2.

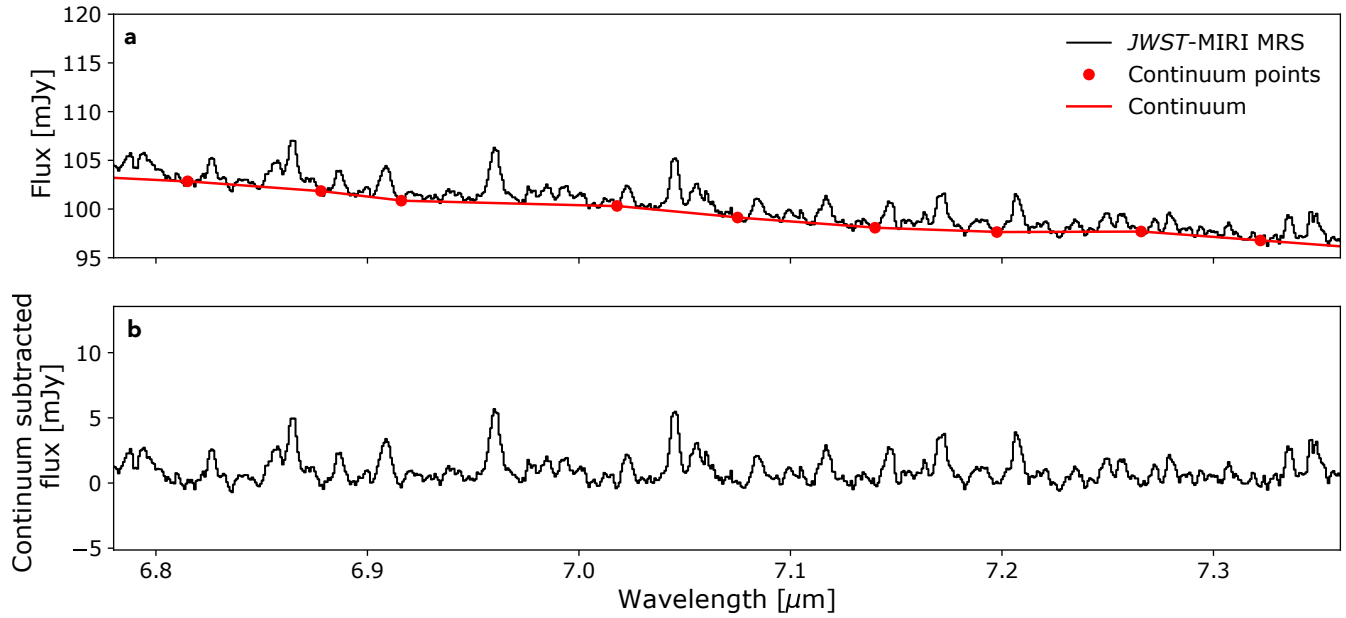
WISE Time-series observations. Extended Data Fig. 6 reports WISE Time-series observations of PDS 70. Observations were executed on February 2-3, and 6, 2010, and on August 1-2, 2010. We note that the source is highly variable and that WISE 4 (W4; $25 \mu\text{m}$) is anticorrelated with WISE 1 (W1; $3.4 \mu\text{m}$) and WISE 2 (W2; $4.6 \mu\text{m}$). Such variability may be "seesaw"-like¹⁹, where changes in the scale height of the inner disk wall shadow the disk material located further out. However, a complete "seesaw" profile is not observed, as at wavelengths shortwards of $8 \mu\text{m}$ the MIRI spectrum lies above the IRS spectrum (Fig. 1). This is not surprising as the wavelength of the "pivot" point (i.e., the wavelength at which a shift in emission is observed) is dependent on the location of the occulting material with respect to the star, the stellar luminosity and the inclination of the system, with highly inclined systems showing a more complete "seesaw" than more face-on systems such as PDS 70 ($i = 51.7 \pm 0.1^\circ$ ¹¹). Interestingly, WISE 3 (W3; $12 \mu\text{m}$) is not anticorrelated with WISE 1 and WISE 2 due to the dominant $10 \mu\text{m}$ silicate band which indeed shows a minor offset compared to the longer wavelengths. It indicates that the material contributing to the $10 \mu\text{m}$ emission is not shadowed. This behaviour is seen if the emission arises from warmer dust closer to the star than the occulting material, or further above the disk midplane.

We also note that the difference in aperture size of *Spitzer*-IRS and MIRI-MRS cannot explain the observed variability. The *Spitzer*-IRS low-resolution spectrograph has a slit width of $3.6''$ for wavelengths shortwards of $14 \mu\text{m}$ and $10.2''$ for wavelengths longwards of $14 \mu\text{m}$. While the maximum aperture of *Spitzer*-IRS at longer wavelengths is larger than that of MIRI-MRS, this is not the case for the shorter wavelengths where the slit widths are similar for both observatories ($\sim 3.6''$ vs $4.0''$). However, a flux offset is also observed in this spectral region. Additionally, in the case that the long wavelength excess would arise from an extended component, a jump in flux level at $14 \mu\text{m}$ - where the aperture size changes - would be present in the IRS data, but it is absent. Furthermore, we note that there is no hint of extended emission in the spatial direction of both IRS and MIRI observations.

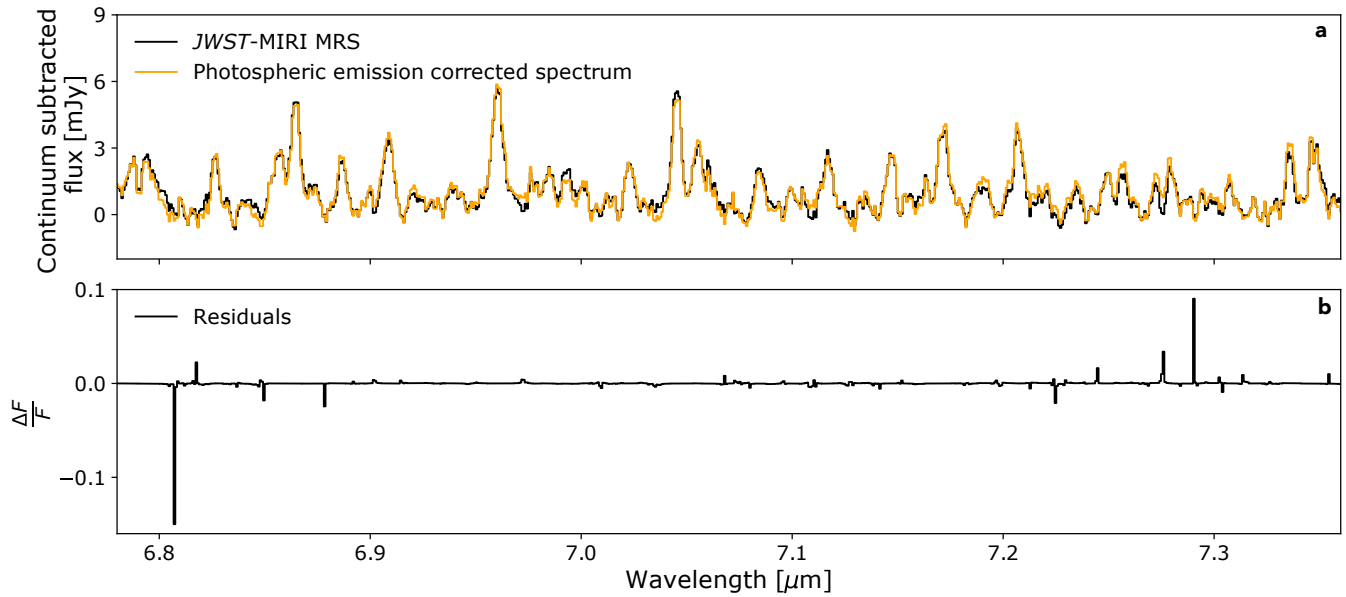
1 Extended Data



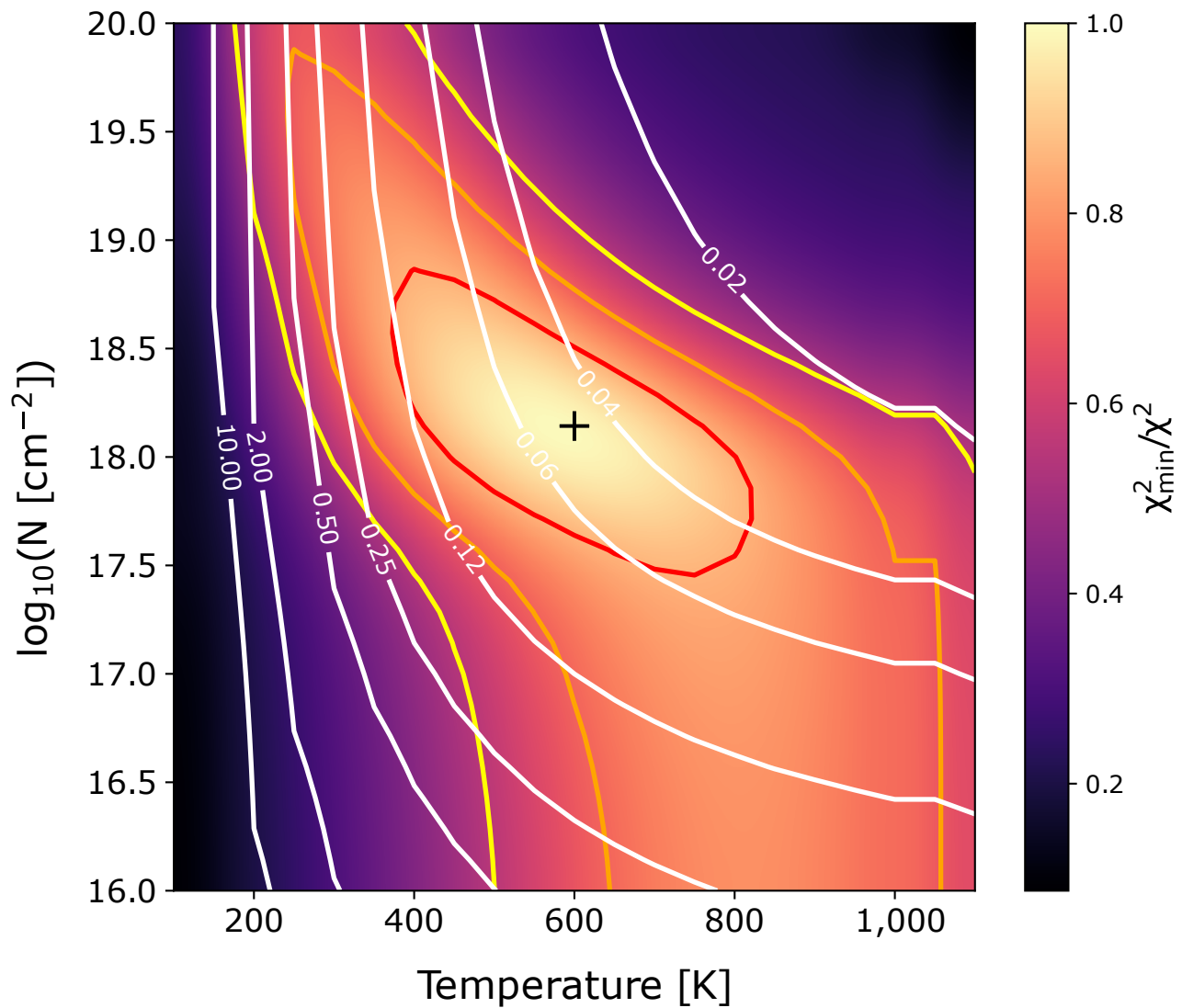
Extended Data Figure 1. The architecture of the PDS 70 system. **a**, Schematic representation of the locations of the inner and outer disk of PDS 70 indicated as teal and blue ellipses. The protoplanets PDS 70 b and PDS 70 c are shown as blue dots. **b**, Main components of the schematic of the system on top of a MIRI-MRS IFU image at 7.0 μm, illustrating the size of the system with respect to the 2.5-FWHM aperture used for spectro-photometric extraction (white circle). The latter linearly increases with wavelength.



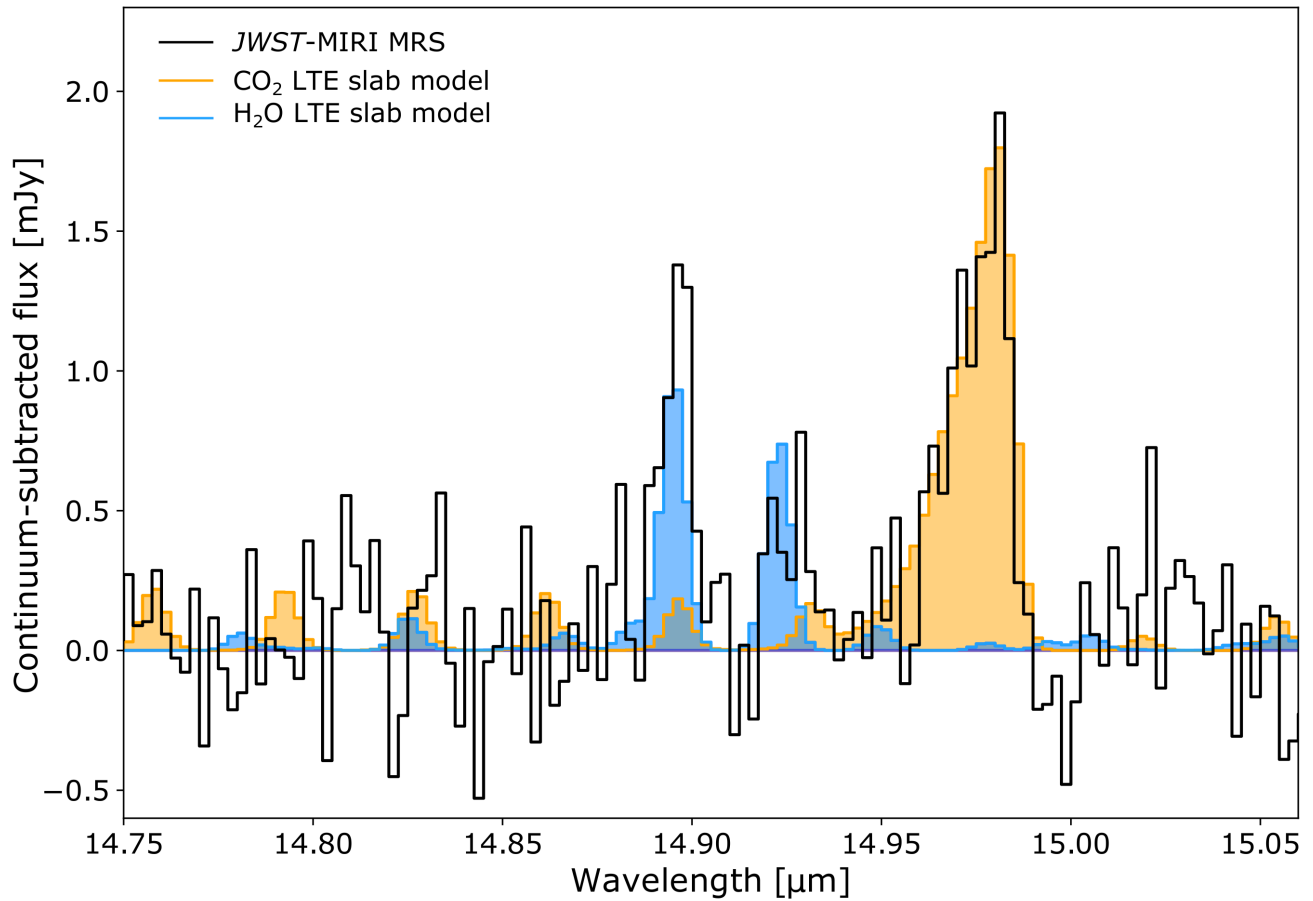
Extended Data Figure 2. Local continuum fit used in the spectrum presented in Figure 3. **a**, The selected continuum points are displayed as red dots and the interpolated continuum is shown as a red line. **b**, The continuum-subtracted spectrum.



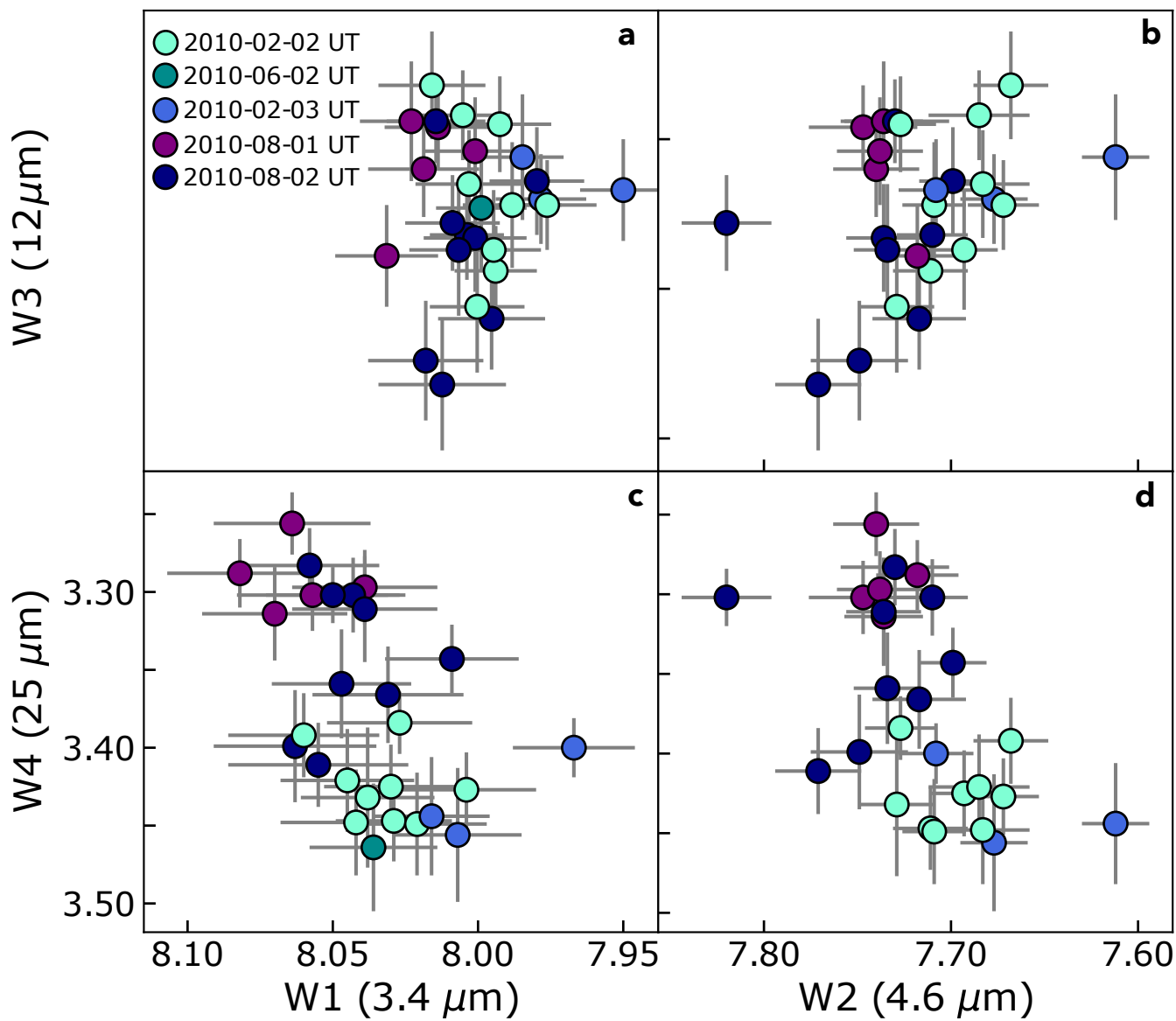
Extended Data Figure 3. **a**, Comparison between the MIRI-MRS spectrum (black) and the spectrum corrected for the stellar photosphere (orange). Both spectra are continuum subtracted. **b**, The residuals show that the contamination from the stellar photosphere is negligible in the observed spectrum.



Extended Data Figure 4. χ^2 map for the fit of the $7 \mu\text{m}$ region of the H_2O bending mode. The best-fit model is represented by a black plus. The 1σ , 2σ , and 3σ confidence intervals are shown in red, orange, and yellow, respectively, for a typical noise level of $\sigma = 0.15$ mJy. The best-fitting emitting radius R for all values of N and T is indicated as white lines. In general, we find a degeneracy between a high T and low N solution, and a low T and high N solution. Within the framework of our LTE slab model, the data indicate mildly optically thick H_2O emission at a temperature of about 600 K.



Extended Data Figure 5. Continuum-subtracted spectrum in the 15 μm region showing the detected Q-branch of CO_2 (orange). The shape of this feature is sensitive to temperature and is well-fitted by an LTE slab model with $T = 200$ K. The strength of this feature can be reproduced with $N(\text{CO}_2) = 1.5 \times 10^{17} \text{ cm}^{-2}$ and $R = 0.1$ AU. However, the aforementioned parameters are degenerate and are used for illustrative purposes only. Rotational lines of H_2O ($J = 14_5 \ 10 - 13_2 \ 11, J = 14_6 \ 9 - 13_3 \ 10; E_u \sim 4300$ K) are also detected and they are reasonably well reproduced by the best-fit model for the 7 μm region (blue). This could indicate that there is no additional reservoir of water at cooler temperature.



Extended Data Figure 6. WISE Time-series photometry of PDS 70. Errorbars represent 1 s.d. **a–b**, WISE 3 (W3) is not anticorrelated with WISE 1 and WISE 2 due to the dominant $10\ \mu\text{m}$ silicate emission which does not vary substantially throughout different epochs. **c–d**, Anticorrelations are observed for WISE 1, WISE 2 and WISE 4 indicating a "seesaw"-like time variability (see Methods for further details).

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